On the rotational behaviour of parent stars of extrasolar planets

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ABSTRACT
We analysed the behaviour of the rotational velocity in the parent stars of extrasolar planets. Projected rotational velocity $v \sin i$ and angular momentum were combined with stellar and planetary parameters, for a unique sample of 147 stars, amounting to 184 extrasolar planets, including 25 multiple systems. Indeed, for the present working sample we considered only stars with planets detected by the radial-velocity procedure. Our analysis shows that the $v \sin i$ distribution of stars with planets along the HR diagram follows the well-established scenario for the rotation of intermediate to low main-sequence stars, with a sudden decline in rotation near 1.2 $M_\odot$. The decline occurs around $T_{\text{eff}} \sim 6000$ K, corresponding to the late-F spectral region. A statistical comparison of the distribution of the rotation of stars with planets and a sample of stars without planets indicates that the $v \sin i$ distribution for these two families of stars is drawn from the same population distribution function. We also found that the angular momentum of extrasolar planet parent stars follows, at least qualitatively, Kraft’s relation $J \propto (M/M_\odot)^{\alpha}$. The stars without detected planets show a clear trend of angular momentum deficit compared to the stars with planets, in particular for masses higher than about 1.25 $M_\odot$. Stars with the largest mass planets tend to have angular momentum comparable to or higher than the Sun.

Key words: planetary systems – stars: rotation.

1 INTRODUCTION
Since the pioneering discovery by Mayor & Queloz (1995) of a Jupiter-mass planet orbiting 51 Peg star, a large number of extrasolar planets have been discovered orbiting stars of different evolutionary stages and masses. As of 2008 July 30, there were 322 planetary-mass companions known to be orbiting mainly solar-type stars, including 35 multiphilet systems. During the past decade these discoveries have inspired intensive studies on the physical properties of the planets and of their parent stars. The discovered planets have masses ranging from 4 Earth masses to 11 Jupiter masses. They can be found at distances of several au or close to the parent star, with orbital periods in the range of a few days to a few years. High eccentricity is a common parameter connecting these planets (e.g. Marcy et al. 2001).

The nearby dwarf stars with giant planets show evidence of moderate metal enrichment compared to the average metallicity of field dwarfs in the solar neighbourhood without detected planets. The dependence of planetary frequency on the metallicity of host dwarf stars has been investigated since they were first detected by precision radial-velocity surveys (Gonzalez 1997; Laughlin & Adams 1997). Different explanations have been proposed for this dependence, such as enhanced giant planet formation by high stellar metallicity (Santos et al. 2000; Santos, Israelian & Mayor 2001; Reid 2002), observational selection effects or pollution by ingested planetary material (Laughlin 2000; Gonzalez et al. 2001; Israelian et al. 2001; Pinsonneault, DePoy & Coffee 2001; Murray & Chaboyer 2002). For instance, based on observationally unbiased stellar samples, the evidence for higher planetary frequency around unpolluted, primordial metal-rich stars has been clearly demonstrated by Santos et al. (2001), and confirmed by Fischer & Valenti (2003) and Santos, Israelian & Mayor (2004), who showed a sharp break in frequency at $[\text{Fe/H}] \sim 0.0$. Further, several studies have investigated a possible link between the orbital period of an extrasolar planet sample and the metallicity of the parent stars. Some authors have argued in favour of this hypothesis (Gonzalez 1998; Queloz 2000; Jones & Sleep 2003), while others (Santos et al. 2001; Laws et al. 2003) found no evidence of its existence. In particular, Santos et al. (2003) concluded that the metallicity distribution of stars with very short-period planets, typically $P_{\text{orb}} < 10$ d, is essentially indistinguishable from the same distribution of stars with longer orbital period planets, whereas Sozzetti (2004) claims there is a possible correlation between stellar metallicity and planet orbital period, where close-in planets, in an orbit of a few days, would more likely be found around metal-rich stars. In contrast to the behaviour of metallicity in dwarf stars with planets, giant stars hosting planets do not show a tendency of being more metal rich (Pasquini et al. 2007). Mazeh, Zucker & Pont (2005) pointed out an intriguing correlation between
the masses and orbital periods of extrasolar planets, consistent with a linear relation, at least for planets with periods shorter than 5 d. Such a result was confirmed more recently by Torres, Winn & Holman (2008). In addition, Gonzalez (2008) confirmed that, near the solar temperature, Li abundances of stars with planets are smaller than those of stars without planets. For a solid review, presenting a census of the main statistical results obtained to date in this domain and a discussion on extrasolar planet general orbital properties, the reader is referred to Udry & Santos (2007).

Among other possible defining characteristics of stars with planets, little attention has been paid to the analysis of their rotational behaviour. More specifically, a comprehensive study on the rotational behaviour of stars harbouring extrasolar planets, in spite of the obvious importance of rotation in stellar and planetary system evolution, has yet to be performed. Actually, the vast majority of recently discovered planetary systems contain solar-type stars and this fact suggests that stars with planets are slow rotators at their respective ages. Indeed, only a preliminary study by Barnes (2001), based on the rotation periods of 35 such stars, focused on this subject, suggesting that planet host stars have normal rotational properties, whereas a more recent study by Gonzalez (2008) points to a scenario where the $v \sin i$ values of stars with planets are smaller than those of stars without planets for stars cooler than 6000 K, nearly the same for stars near 6000 K and much larger for the hottest stars.

This paper presents an unprecedented analysis of the behaviour of the projected rotational velocity $v \sin i$ of stars with planets, using the current sample of such stars available in the literature, in addition to searching for a possible connection between stellar rotation and planetary parameters. An analysis of the angular momentum behaviour of these stars is also performed. To make headway in the study of the rotational behaviour of stars with planets, it is, therefore, essential to conduct a comparative analysis of the rotation and angular momentum behaviour of stars without detected planets. This is one of the major goals of this paper. The paper is organized as follows. In Section 2 we present the characteristics of the working sample and in Section 3 we analyse possible sources of bias that might contribute to producing the features observed. In Section 4 we present our findings, with a brief discussion and finally, conclusions are outlined in Section 5.

2 STELLAR WORKING SAMPLE

This paper is based on the sample of stars with extrasolar planets listed in the comprehensive Extrasolar Planets Encyclopedia, maintained by J. Schneider, as of 2008 July 30, amounting to 322 planetary-mass companions known to be orbiting 272 almost solar-type stars, including 35 multiple planet systems. Nevertheless, in the final working sample we considered only the stars with planets detected by radial-velocity procedure, consisting of 147 stars with projected rotational velocity $v \sin i$ available in the aforementioned encyclopedia, hosting 184 planets, including 25 multiple systems. It should be pointed out that all these stars are listed in the Geneva planetary search survey. With such a criterion, all the stars with transiting planets, selected only from photometric surveys, are purged. This fact automatically eliminates a possible bias favouring fast rotators coming from transiting systems. The main stellar and planetary physical observables used in this paper, namely $v \sin i$, mass and stellar age, as well as the planetary orbital parameters, can be retrieved from Schneider (2009) and references therein. Readers are referred to these references for a discussion of measurement procedures and error analyses of the stellar parameters listed.

For comparative purposes, we took two samples of stars not known to have any planetary-mass companions, one from the Geneva planet search survey composed of 39 F- and G-type dwarf stars (Santos et al. 2004, 2005) and the other from the Anglo-Australian planet survey composed of 85 dwarf stars (Bond et al. 2008). However, we should be cautious about these samples, which derive from a list of stars that are surveyed for planets, but for which none has yet been found. Certainly, this does not mean that such stars have no planetary companions whatsoever. For instance, they might host planets with very low mass and/or long orbital period that are more difficult to detect with radial-velocity surveys.

Because projected rotational velocity $v \sin i$ is perhaps the most sensitive parameter in this paper, let us briefly discuss the quality and precision of the data obtained. For stars with planets, $v \sin i$ measurements were taken from Valenti & Fischer (2005) (91 stars), Holmberg, Nordström & Andersen (2007) (26 stars), Marcy & Chen (1992) (four stars), da Silva et al. (2006, 2007) (four stars), Fischer et al. (2007) (three stars), Bakos et al. (2007a,b) (two stars), Lovis et al. (2005) (two stars), Santos et al. (2007) (two stars), Tamuz et al. (2008) (two stars) and one star each from Bonfils et al. (2007), Burke et al. (2007), Ge et al. (2006), Laughlin et al. (2005), Mayor et al. (2009), McCullough et al. (2006), Naef et al. (2007), O’Donovan et al. (2006), Winn et al. (2007), Johnson et al. (2006) and Wright et al. (2007). Measurements from Holmberg et al. (2007) were computed using the CORAVEL cross-correlation method, with typical errors of about 1 km s$^{-1}$, whereas those from Valenti & Fischer (2005) were computed using the spectral synthesis procedure, with errors of about 0.5 km s$^{-1}$. Fig. 1 presents the error distribution in $v \sin i$, ranging from about 0.5 to 2.5 km s$^{-1}$, with a mean error of 0.6 ± 0.3 km s$^{-1}$.

For both stellar samples, the stellar effective temperature was computed from the $T_{\text{eff}} = (B - V) \text{ calibration given by Flower}$ (1996), whereas the stellar radius was estimated according to Lang (1980). Stellar luminosity $L$ was estimated following the same procedure employed by do Nascimento et al. (2000), using the parallaxes and the $m_V$ magnitudes given by Hipparcos to derive the
intrinsic absolute magnitudes $m_v$. Stellar ages, all computed on the basis of isochrone interpolation (e.g. Saffe & Gomez 2005; da Silva et al. 2007; Holmberg et al. 2007), were taken from Schneider (2009) for stars with planets, and from Valenti & Fischer (2005) for stars without planets, whereas masses, all determined spectroscopically, were taken from Schneider (2009) and Valenti & Fischer (2005), respectively, for stars with and without planets.

We also analysed the behaviour of stellar angular momentum, which was computed by assuming stars to be a solid and uniform density sphere. For the calculation of this parameter we used the relation for the mean stellar angular momentum $J(M)$, given by

$$ J(M) = \frac{v \sin i}{R(M)} I(M), $$

where $I(M) = (2/5)MR^2$ is the solid body moment of inertia for a sphere, $v \sin i$ is the projected stellar rotational velocity and $R(M)$ the radius of the star.

### 2.1 Metallicity bias and $v \sin i$ of stars with planets

As stated by Gonzalez (2003), three main sources of bias have been identified as being potentially important to the study of the physical properties of stars with planets:

(i) Biases linked to specific determination of the metal content of the host stars. As pointed out by Laws et al. (2003), spectroscopic determination of [Fe/H] seems to be more reliable than photometric methods.

(ii) Biases in statistical analyses because of the sample limitation of stars with [Fe/H] < 0.0. (iii) Biases linked to intrinsic metal line behaviour.

The fundamental method used to analyse the influence of these biases comes from Santos et al. (2003) and Fischer & Valenti (2003). By calculating the median velocity error as a function of metallicity for the stars in their planet surveys, these authors found a velocity degradation of up to 50 per cent for the lowest metallicity stars ([Fe/H] ~ 0.5). Radial-velocity surveys currently attain typical single measurement precision $\sigma_{RV} \sim 3-5$ m s$^{-1}$, indicating that planets would be easily detected even with $\sigma_{RV} \sim 5-8$ m s$^{-1}$. One can therefore conclude that radial-velocity precision degradation for metal-poor stars is not the major cause of any correlation. To explain the fact that the lower occurrence rate of stars with planets around metal-poor stars had little influence on our results, we divided the working sample of 147 stars into three bins: (i) $-0.71 \leq$ [Fe/H] < 0.00, slightly underabundant in [Fe/H]; (ii) 0.00 \leq [Fe/H] < 0.21, around solar abundance and (iii) 0.21 \leq [Fe/H] < 0.50, slightly overabundant in [Fe/H]. For this analysis we used only spectroscopic determinations of [Fe/H]. The distributions for the three different bins are displayed in Fig. 2, showing no clear trend for $v \sin i$ in the metallicity ranges considered.

### 3 RESULTS AND DISCUSSION

The aim of this pioneering study is to initiate a fruitful discussion to make the community aware of the fundamental role of rotation in our understanding of extrasolar planets and their relation with central stars. To this end, we dedicated most of our efforts to identifying qualitative trends between rotation and different stellar and planetary parameters.

#### 3.1 The rotation of stars with planets

The distributions of the projected rotational velocity $v\sin i$ in the HR diagram for our working samples of 147 stars with planets and 85 stars without detected planets is displayed, respectively, in the upper and lower panels of Fig. 3. The Sun is also indicated for comparative purposes. Evolutionary tracks, with solar metallicity, follow the procedure used by do Nascimento et al. (2000). The first interesting feature emerging from this figure is the fact that the present sample of main-sequence stars with planets, once displayed in the HR diagram, tends to follow the well-established scenario for the rotation of intermediate to low main-sequence stars, with a sudden decline in $v \sin i$ values near 1.2 M$_\odot$. Such a scenario was first pointed out in the remarkable paper of Kraft (1967) on the strong observational basis emerging from $v \sin i$ measurements for late-F and early-G dwarfs in nearby young clusters and in the field. This decline occurs around $T_{\text{eff}} \sim 6000$ K, corresponding to the late-F spectral region (e.g. Soderblom 1983). Indeed, a detailed inspection of the aforementioned figure shows that the rotational behaviour of main-sequence stars with planets parallels that observed in the milestone HR diagram presented by Kraft (1967), where stars with enhanced rotation are mostly those with effective temperature higher than about 5800 K and luminosity above the solar value, whereas slow rotators are mostly cooler stars, with luminosity lower than the solar value. The fast and moderate transition to slow rotators is explained as resulting from the presence of a convective envelope in late-type stars that slows them down from an initially rapid rotation rate through magnetic breaking, associated to the fact that these stars are on average older than the early-type stars and have undergone a longer spin-down period. In spite of the smaller number of stars without detected planets, relative to the sample of stars with planets, a comparative inspection of panels displayed in Fig. 3 shows a clear trend for similar distributions of the projected rotational velocity $v \sin i$ for both families of stars, in the region of the HR diagram defined by $0.6 \gtrsim \log(L/L_\odot) \gtrsim -0.2$ and $3.8 \gtrsim \log T_{\text{eff}} (K) \gtrsim 3.7$.

In spite of the scenario observed in the distribution behaviour of $v \sin i$ in planet host stars along the HR diagram, while following the scenario discovered by Kraft (1967), one important question
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Figure 3. The distribution of rotational velocity $v \sin i$ for 147 parent stars of extrasolar planets (upper panel) and for 85 stars without detected planets from Bond et al. (2008) (lower panel). Evolutionary tracks for solar metallicity followed the procedure used by do Nascimento et al. (2000). The sun is also represented for comparative purposes.

emerges at this point: are the rotational properties of planet host stars normal compared to the stars without detected planets? Although objective comparisons between the parameters of stars with planets and those of stars without planets have proven to be difficult, this paper offers the possibility of a statistical comparison between the rotation distributions in these two families of stars. Indeed, the selection criteria of the present radial-velocity planet search surveys include stars exhibiting sufficiently sharp-lined spectra, to detect the small Doppler shifts induced by planetary-mass companions and, in this context, both samples defined in the previous section exclude stars with very large $v \sin i$ values, exhibiting only slow to moderate rotators.

In order to verify if the present data sets are significantly different from one another, we performed a Kolmogorov–Smirnov (KS) test (Press et al. 1992), which calculates the probability that two distributions are derived from the same parent distribution. According to the KS test, zero probability means the distributions are dissimilar, whereas unit probability means they are the same. During the present study we applied the Kuipers KS statistic (e.g. Jetsu & Pelt 1996; Paltani 2004), hereafter the KS test, which, in contrast to the standard KS statistic, is invariant under an origin shift for distributions, offering the same weight for all the points of a distribution.

As a first step, we performed KS analyses, taking into account all the main-sequence stars listed in our working sample, namely 147 stars with planets and 85 stars without detected planets. Fig. 4 shows the cumulative functions for both $v \sin i$ distributions. The probability value of 0.209 obtained on the KS test is consistent with the two distributions being drawn from the same population. In addition, we performed the KS analyses by comparing groups of stars that are in the same region of the HR diagram. Samples of dwarf stars with and without planets were grouped first by spectral types, namely F- and G-type stars, and then by effective temperature, namely stars cooler and hotter than 6000 K. Now taking into account the sample of 118 F- and G-type main-sequence stars with planets, defined in this paper, and 82 F- and G-type stars without planets, listed by Bond et al. (2008), a probability value of the KS test of 0.299 is obtained, reinforcing the previous scenario with the two distributions of $v \sin i$ being drawn from the same parent population. In addition, if we compare similar samples of F- and G-type dwarf stars from the Geneva survey, a probability value of 0.947 is obtained, indicating that $v \sin i$ distributions are drawn from the same parent distribution, at a 93 per cent confidence level.

The KS test applied for main-sequence stars with planets segregated by effective temperature, namely cooler and hotter than 6000 K, gives a probability of $4 \times 10^{-6}$, pointing towards a scenario where the distributions of $v \sin i$ for stars with planets, hotter and cooler than 6000 K are in fact not drawn from the same population distribution function. Such a result reinforces the scenario observed in Fig. 3, with a clear decline in $v \sin i$ of dwarf stars with planets around $T_{\text{eff}} \sim 6000$ K.

Figure 4. Comparison of the distribution of rotational velocity $v \sin i$ of main-sequence stars with and without planets. Solid and dashed curves stand for the cumulative distribution of $v \sin i$ of stars with and without planets, respectively.
To further explore a possible relationship between the rotation of stars and the presence of planets, we also applied the KS test for $v \sin i$ distribution of stars with one, two or more planets, looking for a possible reflex of the number of planetary companions in the rotation of their parent stars. Fig. 5 shows the cumulative $v \sin i$ functions for stars with one and two or more planets. In the present case, a probability value of 0.574 was obtained, indicating that the $v \sin i$ distributions for stars with one and two or more planets are consistent with samples being drawn from the same population distribution. However, these results should be interpreted cautiously, since they may be a result of biases associated with $v \sin i$ measurement uncertainties or selection effects in the definition of the two stellar samples, including stellar ages and masses. Accordingly, we performed an additional KS test to check the statistical nature of the distributions of effective temperature and age of the stars with and without detected planets. Fig. 6 shows the cumulative functions for the distributions of effective temperature (upper panel) and age (lower panel), for both the samples of F- and G-type main-sequence stars. Probability values of 0.268 and 0.236 were obtained with the KS test for effective temperature and age, respectively, which is consistent with the present samples of stars with and without planets coming from populations with identical distributions in the HR diagram. The same result was obtained with the KS test of the distributions of luminosity for the aforementioned samples, with a probability value of 0.121. It must be remembered that stellar ages from isochrones have an uncertainty that varies dramatically with location in the HR diagram, a fact that may considerably compromise the KS test for ages, in spite of the fact that we are considering only F- and G-type main-sequence stars. Nevertheless, the most relevant point in our comparative analyses is the KS test for $T_{\text{eff}}$ and luminosity, which presents realistic uncertainties.

### 3.2 The angular momentum of stars with planets

One notable characteristic of the Solar system, with direct implications for its formation process, is the fact that most of its angular momentum is in planetary orbital motion. For instance, total orbital angular momentum of the Solar system, largely due to Jupiter, is at least two orders of magnitude larger than the spin angular momentum of the Sun. By extending the study of angular momentum behaviour to other stars with planets, we can examine how different or similar their spins are compared to the Sun. In the following, we speculate very briefly on the angular momentum characteristics of the parent stars of extrasolar planets, by comparing the...
angular momentum of stars with and without detected planets. Further, we analyse the angular distribution for the two samples of stars in the context of Kraft’s well-known law $J \propto (M/M_\odot)^\alpha$ (Kraft 1967; Kawaler 1987).

For this purpose, we estimated the angular momentum for the F- and G-type dwarf stars of the present working samples, stars with and without planets, according to the recipe described in Section 2.1. Let us recall that, for the computation of the angular momentum we used masses, determined spectroscopically by Schneider (2009) and Valenti & Fischer (2005), for stars with and without planets, respectively. Indeed, as pointed out by Valenti & Fischer (2005), their spectroscopically computed masses have a median fractional precision of 15 per cent and are systematically 10 per cent higher than masses from interpolating isochrones. Such a fact explains the presence of a dozen stars in our analyses with $M \geq 1.25 \text{M}_\odot$, in contrast to the distribution of stars without planets in the HR diagram displayed in the previous section. Fig. 7 shows the distribution of the angular momentum of the main-sequence F- and G-type stars that make up our working sample, as a function of stellar mass (in solar mass unity). In this figure, open and solid symbols stand for stars with and without planets, respectively. The solid line represents the best fit of Kraft’s law $J \propto (M/M_\odot)^\alpha$ (Kraft 1967; Kawaler 1987).

As a check of the patterns observed in Fig. 7, we applied the procedure proposed by Gonzalez (2008) for comparison of stellar properties of stars with and without planets, which estimate a measure of the proximity of two stars in $[T_{\text{eff}}, \log g, [\text{Fe/H}], M_i]$ space. Panel (a) of Fig. 8 shows the angular momentum differences $(\log J_{\text{plan}} - \log J_{\text{comp}})$ between stars with planets $(J_{\text{plan}})$ and comparison stars $(J_{\text{comp}})$ represented in Fig. 7. Panel (a) shows the differences between stars with planets and the most similar comparison stars, from the stellar samples represented in Fig. 7, whereas panel (b) corresponds to stars with and without planets from the Geneva planet search survey.

We turn now to an analysis of the role of planetary mass on the angular momentum $J(M)$ of the parent stars. Fig. 9 shows the stellar angular momentum distribution of 118 F- and G-type
dwarf stars with planets, now separated into three mass intervals of their orbiting planets, namely $2.0M_{\text{Jup}} < M_{\text{pl}} \leq 5.0M_{\text{Jup}}$, $5.0M_{\text{Jup}} < M_{\text{pl}} \leq 8.0M_{\text{Jup}}$, and $8.0M_{\text{Jup}} < M_{\text{pl}}$, represented in panels (a), (b) and (c), respectively. A close inspection of these panels shows that stars hosting more massive planets tend to have the highest angular momentum, whereas stars with the least massive planets tend to have the lowest. Interestingly, stars hosting planets with the largest mass tend to have angular momentum larger than the Sun.

4 CONCLUSIONS

This paper conducts an unprecedented analysis of the rotational velocity behaviour of the parent stars of extrasolar planets. We combined the projected rotational velocity $v \sin i$ of stars with planets with stellar and planetary parameters, for a unique sample of 147 stars, amounting to 184 extrasolar planets discovered by the radial-velocity procedure, including 25 multiple systems, searching for trends or anomalies in their relations. In addition, for comparative purposes, we used a sample of stars without detected planets composed of 85 dwarf stars from the Anglo-Australian planet survey. Different KS tests, performed with effective temperature, luminosity and age, indicate that both samples of stars come from populations with similar distributions in the HR diagram.

First, we placed stars in the HR diagram, from which one observes that stars with planets follow the well-established scenario for the rotation of intermediate to low main-sequence stars, first pointed out by Kraft (1967), with a sudden decline in rotation near $2M_\odot$. Stars with planets exhibiting enhanced rotation are mostly those with effective temperature higher than about 6000 K and luminosity higher than the solar value, whereas slow rotators are mostly the coolest stars, with luminosity lower than the solar value. This transition, fast rotators and moderate to slow rotators, is explained as resulting from the presence of a convective envelope in late-type stars. This convective envelope slows down the stars from an initially rapid rotation rate through magnetic breaking, associated to the fact that these stars, on average, are older than the early-type stars and have undergone a longer spin-down period. Qualitatively, stars without detected planets appear to follow the same scenario. The distributions of $v \sin i$ for stars with and without detected planets suggest that they are being drawn from the same parent population.

We also analysed the stellar angular momentum behaviour as a function of stellar mass, in which one observes that stars with planets follow, at least qualitatively, Kraft’s law $J \propto (M/M_\odot)^{\alpha}$ (Kraft 1967; Kawaler 1987). Further, we compared the distribution of angular momentum $J(M)$ in parent stars of extrasolar planets versus that of stars without detected planets. The latter shows a clear trend of being in angular momentum deficit compared to the stars with planets. Furthermore, we observed that the vast majority of stars without detected planets are also in angular momentum deficit compared to the Sun. If these tentative conclusions are confirmed, it would imply that the angular momentum distribution in these two families of stars followed different histories, that is, stars with detected planets following the Solar system history. Finally, we analysed the distribution of the angular momentum of stars with planets as a function of planetary mass. Here, stars with more massive planets tend to have the highest angular momentum in relation to the solar value.

As we underscored previously, the sample of stars hosting planets is subject to a number of subtle biases, some showing the effect of masking the rotational behaviour of such stars, including the relationship between stellar rotation and planetary orbital parameters. Perhaps the major bias in the present study arises from the difficulty of detecting planets around fast rotators, since, to date, the main planet detection process, namely the Doppler method, cannot account for stellar spectra with broad lines. A large $v \sin i$ value results in broad lines and is a clear source of noise in Doppler measurements (Wright 2005). Again, planetary mass may also represent an additional observational bias in the study of the link between rotation and stellar and planetary parameters, in particular because the actual masses of the vast majority of extrasolar planets are unknown. In addition, as already highlighted, all studies on plan-etary mass show a strong observational selection effect stemming from the difficulty of detecting low-mass long-period planets, due to their small radial-velocity amplitudes. Despite these limitations, the present work points to some interesting trends, which deserve a more detailed analysis to draw firmer conclusions, including a much larger detection sample, particularly of stars with low-mass planets and large orbital periods.
The most interesting finding in this paper is undoubtedly the discrepancy in the angular momentum behaviour of stars with and without planets, with the majority of stars without detected planet exhibiting a deficit in angular momentum compared to hosting planet stars. If stars with as yet undetected planets are also hosting planetary systems, their planets should have different masses and orbital parameters from those already detected. Such an aspect may have a strong impact on the angular momentum behaviour and structure of these stars. For instance, one can expect that the total angular momentum of these as yet undetected planetary systems is more distributed among the planets, an aspect reflected by a deficit in the angular momentum of the star compared to present-day stars with detected planets.

The study of the rotation behaviour in a larger sample of stars with planets, detected by the transit procedure, may help us clarify some of the trends observed in this paper involving planetary mass. Infrared modulation in planets orbiting stars with enhanced rotation could also reveal a signature of pseudo-synchronized planetary rotational periods. Asteroseismological studies of stars with and without planets may also help us elucidate the discrepancy in the angular momentum of these two families of stars, revealing peculiarities in the inner stellar structure with direct impact on the angular momentum evolution, such as the degree of differential rotation. The different patterns observed during this work could also have been influenced by important biases associated to the small sample of stars without planets and different precisions in the $v \sin i$ measurements of both stellar samples. An essential first step is to determine the $v \sin i$ measurements of both stellar samples using the same procedure.

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